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Useful constants

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|--------------------------------|------------------|------------------------|---|
| Speed of light | c | $2.998 \cdot 10^{10}$ | cm sec ⁻¹ |
| Planck's constant | h | $6.626 \cdot 10^{-27}$ | erg s |
| Rationalized Planck's constant | $\hbar = h/2\pi$ | $1.055 \cdot 10^{-27}$ | erg s |
| Boltzmann's constat | k_B | $1.381 \cdot 10^{-16}$ | erg °K ⁻¹ |
| Electron charge | e | $4.803 \cdot 10^{-10}$ | esu |
| Electron rest mass | m_e | $9.110 \cdot 10^{-28}$ | g |
| Gravitational Constant | G | $6.673 \cdot 10^{-8}$ | dyne cm ² g ⁻² |
| Avogadro's number | N_A | $6.022 \cdot 10^{23}$ | mole ⁻¹ |
| Atomic mass unit | a.m.u. | $1.661 \cdot 10^{-24}$ | g |
| Gas constant | R | $8.314 \cdot 10^7$ | erg °K ⁻¹ mole ⁻¹ |
| Stephan-Boltzmann constant | σ | $5.670 \cdot 10^{-5}$ | erg cm ⁻² s ⁻¹ °K ⁻⁴ |
| Radiation constant | a | $7.564 \cdot 10^{-15}$ | erg cm ⁻³ °K ⁻⁴ |
| Thomson cross section | σ_T | $6.656 \cdot 10^{-25}$ | cm ² |
| Astronomical unit | a.u. | $1.496 \cdot 10^{13}$ | cm |
| Parsec | pc | $3.086 \cdot 10^{18}$ | cm |
| Light year | l.y. | $9.460 \cdot 10^{17}$ | cm |
| Solar mass | M_\odot | $1.989 \cdot 10^{33}$ | g |
| Solar radius | R_\odot | $6.960 \cdot 10^{10}$ | cm |
| Solar luminosity | L_\odot | $3.826 \cdot 10^{33}$ | erg s ⁻¹ |
| Electron volt | eV | $1.602 \cdot 10^{-12}$ | erg |

Black body Radiation

Flux, energy density and pressure of a black body radiation

$$F = \sigma T^4 \quad \epsilon = aT^4 \quad P = \frac{1}{3}aT^4 \quad (1)$$

Potential Energy and Virial theorem

Gravitational potential energy of spherically symmetric mass M :

$$E_{grav} = -G \int_0^M \frac{mdm}{r} \quad (2)$$

Virial theorem (spherical system in hydrostatic equilibrium with vanishing surface pressure) :

$$3 \int_0^M \frac{P}{\rho} dm + E_g = 0 \quad (3)$$

The thermal energy density (*i.e.* per unit mass) is

$$\epsilon_{th} = \frac{P}{(\gamma - 1)\rho}, \quad (4)$$

where $\gamma \equiv \frac{C_p}{C_V} = \frac{n_f + 2}{n_f}$ and n_f is the number of degrees of freedom per particle. Also, recall that for adiabatic processes $P \propto \rho^\gamma$.

Ideal Gas (mono-atomic non-degenerate) Pressure

$$P_{gas} = \frac{\rho}{\mu m_p} k_B T \quad (5)$$

Molecular Weight

For a fully ionized plasma made of H (mass fraction X), He (mass fraction Y) and heavier elements (mass fraction $Z = 1 - X - Y$):

$$\frac{1}{\mu} \approx 2X + \frac{3}{4}Y + \underbrace{\left\langle \frac{(1+Z)}{A} \right\rangle}_{\approx 1/2} Z \quad (6)$$

Cold Electron Degeneracy Pressure

The Fermi momentum is $p_F = \left(\frac{3n_e h^3}{8\pi}\right)^{1/3}$ and the Fermi energy is $\epsilon_F = \frac{p_F^2}{2m}$.

Non-relativistic electrons:

$$P_{e,nr} = \left[\frac{1}{20} \left(\frac{3}{\pi}\right)^{2/3} \frac{h^2}{m_e m_p^{5/3}} \left(\frac{Z}{A}\right)^{5/3} \right] \rho^{5/3} \quad (7)$$

Extremely relativistic electrons:

$$P_{e,r} = \left[\frac{1}{8} \frac{3}{\pi} \frac{hc}{m_p^{4/3}} \left(\frac{Z}{A}\right)^{4/3} \right] \rho^{4/3} \quad (8)$$

Equations for stellar structure

$$\frac{dm}{dr} = 4\pi r^2 \rho \quad (9)$$

$$\frac{dP}{dr} = -\rho \frac{Gm(r)}{r^2} \quad (10)$$

$$\frac{dT}{dr} = -\frac{3k\rho}{16\pi a c r^2 T^3} L \quad (11)$$

$$\frac{dL}{dr} = 4\pi r^2 \rho \epsilon \quad (12)$$

The opacity (absorption coefficient per unit mass; κ) is approximated at $T > 10^5 K$ by : $\kappa = \kappa_0 \rho T^{-3.5}$ (Kramer opacity)

The energy production rate (per unit mass) ϵ is approximated by $\epsilon = \epsilon_0 \rho^m T^\eta$. For pp burning $m = 1$, $\eta \approx 5$. For CNO burning $m = 1$, $\eta \approx 20$. For 3α burning $m = 2$, $\eta \approx 40$.

Convective energy transfer

Condition for convection:

$$\frac{d \ln P}{d \ln \rho} > \gamma \quad (13)$$

When the energy transfer is dominated by convection:

$$\frac{dT}{dr} = \left(\frac{dT}{dr} \right)_{adiab} = \frac{\gamma - 1}{\gamma} \frac{dP}{dr} \quad (14)$$

Nuclear Reactions

Gamov energy:

$$E_G = \left(\pi \alpha Z_A Z_B \right)^2 2mc^2 \quad (15)$$

Gamov factor:

$$g(E) = \exp \left[- \sqrt{\frac{E_G}{E}} \right] \quad (16)$$

Galaxies & Clusters

Rotation curve:

$$V_{circ}^2 \equiv \frac{Gm(r)}{r} \quad (17)$$

Virial theorem for stellar systems:

$$\frac{\ddot{I}}{2} = E_g + 2E_k \quad (18)$$

(where I is the moment of inertia, E_g and E_k are the gravitational and kinetic energies, respectively).

Cosmology

Redshift: $z \equiv \frac{\lambda_{obs} - \lambda_{em}}{\lambda_{em}}$

Hubble's constant: $H_0 = 100hkm/s/Mpc$

Robertson-Walker Metric:

$$ds^2 = (cdt)^2 - R(t)^2 \left[\frac{dr^2}{1 - Kr^2} + r^2(d\theta^2 + \sin^2 \theta d\phi^2) \right] \quad (19)$$

where $K = 0, \pm 1$.

Einstein Field Equations for the RW metrics:

$$\ddot{R} = -\frac{4\pi G}{3} \left(\rho + 3\frac{p}{c^2} \right) R \quad (20)$$

$$R\ddot{R} + 2\dot{R}^2 = 4\pi G \left(\rho - \frac{P}{c^2} \right) R^2 - 2Kc^2 \quad (21)$$

Friedmann Equations:

$$\dot{R}^2 = \frac{8}{3}\pi G\rho R^2 - Kc^2 \quad (22)$$

Definition: $a(t) \equiv \frac{R(t)}{R(t_0)}$

Cosmological parameters:

$$H_0 = \left(\frac{\dot{R}}{R} \right)_0 \quad (23)$$

$$\Omega_0 = \left(\frac{\rho}{\rho_c} \right)_0 \quad (24)$$

$$\rho_{c,0} = \frac{3H_0^2}{8\pi G} \quad (25)$$

$$q_0 = -\left(\frac{\ddot{R}R}{\dot{R}^2} \right)_0 \quad (26)$$

'Cosmological' equation of state:

$$p = w\rho c^2 \quad (27)$$

Cosmological constant:

$$\rho_\Lambda = \frac{\Lambda c^2}{8\pi G} \quad (28)$$

$$P_\Lambda = -\frac{\Lambda c^4}{8\pi G} \quad (29)$$

Black Body Radiation:

$$\epsilon(\nu) = \frac{8\pi h}{c^3} \frac{\nu^3 d\nu}{\exp(\frac{h\nu}{k_B T}) - 1} \quad (30)$$

$$n_\gamma = \frac{2.4k_B^3}{\pi^2 \hbar^3 c^3} T^3 \quad (31)$$

$$(h\nu)_{mean} \approx 2.7k_B T \quad (32)$$

Luminosity distance:

$$d_L = \left(\frac{L}{4\pi F}\right)^{1/2} \quad (33)$$

Proper distance (at time t_0) to a distance located at coordinate r :

$$d_p(t_0) = R(t_0) \int_0^r \frac{d\tilde{r}}{\sqrt{1 - K\tilde{r}^2}} \quad (34)$$

Lookback time:

$$t_0 - t(z) = \frac{1}{H_0} \left[z - \left(1 + \frac{q_0}{2}\right) z^2 + O[z^3] \right] \quad (35)$$

The $r - z$ relation:

$$R_0 r(z) = c \frac{1}{H_0} \left[z - \left(\frac{1 + q_0}{2}\right) z^2 + O[z^3] \right] \quad (36)$$